

Matter bounce cosmology with a generalized single field: non-Gaussianity and an extended no-go theorem

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Abstract. We extend the matter bounce scenario to a more general theory in which the background dynamics and cosmological perturbations are generated by a k -essence scalar field with an arbitrary sound speed. When the sound speed is small, the curvature perturbation is enhanced, and the tensor-to-scalar ratio, which is excessively large in the original model, can be sufficiently suppressed to be consistent with observational bounds. Then, we study the primordial three-point correlation function generated during the matter-dominated contraction stage and find that it only depends on the sound speed parameter. Similar to the canonical case, the shape of the bispectrum is mainly dominated by a local form, though for some specific sound speed values a new shape emerges and the scaling behaviour changes. Meanwhile, a small sound speed also results in a large amplitude of non-Gaussianities, which is disfavored by current observations. As a result, it does not seem possible to suppress the tensor-to-scalar ratio without amplifying the production of non-Gaussianities beyond current observational constraints (and vice versa). This suggests an extension of the previously conjectured no-go theorem in single field nonsingular matter bounce cosmologies, which rules out a large class of models. However, the non-Gaussianity results remain as a distinguishable signature of matter bounce cosmology and have the potential to be detected by observations in the near future.

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1 Introduction

Matter bounce cosmology [1] is a very early universe structure formation scenario alternative to the paradigm of inflationary cosmology (see, e.g., [2] for a review of inflation, its problems and its alternatives). The idea is that quantum fluctuations exit the Hubble radius in a matter-dominated contracting phase before the Big Bang, which generates a scale-invariant power spectrum of curvature perturbations [3, 4]. The contracting phase is then followed by a bounce and the standard phases of hot Big Bang cosmology. This construction solves the usual problems of standard Big Bang cosmology such as the horizon and flatness problems, but in addition, it is free of the trans-Planckian corrections that plague inflationary cosmology [5], and one can naturally avoid reaching a singularity at the time of the Big Bang (contrary to standard¹ inflation [7, 8]) under the assumption that new physics appears at high energy scales [1, 2]. Nonsingular bounces can be constructed in various ways using matter violating the Null Energy Condition (NEC), with a modified gravity action, or within a quantum theory of gravity (see the reviews [1, 9–12] and references therein).

¹The singularity before inflation could be avoided with, for example, bounce inflation (e.g., [6]).

A typical way of constructing a nonsingular matter bounce cosmology is to assume the existence of a new scalar field. With a canonical Lagrangian, the oscillation of the scalar field can drive a matter-dominated contracting phase when the ratio of the pressure to the energy density averages zero. As the energy scale of the universe increases, new terms can appear in the Lagrangian that violate the NEC and drive a nonsingular bounce. For example, using a Galileon scalar field [13] (or equivalently, in Horndeski theory [14]), one can construct a stable NEC violating nonsingular bounce [15–20] that may be free of ghost and gradient instabilities [21, 22] (see, however, the difficulties in doing so as pointed out by [23–26]).

To distinguish the matter bounce scenario from inflation observationally, studying primordial non-Gaussianities is a useful tool². In the case of inflation, after the calculation of the bispectra generated in single field slow-roll models [29], there have been many studies in the past decade trying to extend the simplest result, which largely enriched the phenomenology of nonlinear perturbations (see [30, 31] for reviews). In particular, one important progress has been to generalize the canonical inflaton to a k -essence scalar field [32, 33], such as k -inflation [34, 35] and DBI models [36, 37], which are collectively known as general single field inflation [38]. In these models, due to the effects of a small sound speed, the amplitude of the bispectrum is enhanced and interesting shapes emerge [30, 31, 38–41]. In a matter-dominated contracting phase, the calculation of the bispectrum has only been done by [42] for the original matter bounce model with a canonical scalar field. A natural extension is thus to consider a k -essence scalar field³ similarly to what has been done in inflationary cosmology, especially since the appearance of a noncanonical field is quite common in the literature of nonsingular bouncing cosmology in order to violate the NEC as explained above. Because the perturbations behave differently in matter bounce cosmology compared to inflation, in particular due to the growth of curvature perturbations on super-Hubble scales during the matter-dominated contracting phase, the canonical matter bounce yields non-Gaussianities with negative sign and order one amplitude, which differs from the results in canonical single field inflation. It would be interesting to explore how these non-Gaussianity results change when one generalizes the original matter bounce scenario to be based on a k -essence scalar field.

Besides non-Gaussianity, another interesting observable for very early universe models is the tensor-to-scalar ratio r . In the original matter bounce scenario, this ratio is predicted to be very large [49, 50]. Indeed, the scalar and tensor power spectra share the same amplitude, and accordingly, the tensor-to-scalar ratio is naturally of order unity [51]. This is well beyond the current observational bound from the Cosmic Microwave Background (CMB), which states that $r < 0.07$ at 95% confidence [52].

A resolution to this problem is to allow for the growth of curvature perturbations during the bounce phase, which suppresses the tensor-to-scalar ratio. However, curvature perturbations tend to remain constant through the bounce phase on super-Hubble scales [20, 53]. In fact, amplification can only be achieved under some tuning of the parameters, and the overall growth is still limited⁴ [51]. Yet, if the scalar modes are amplified, another problem follows

²Another observable quantity, besides non-Gaussianities, that would allow one to differentiate between inflation and the matter bounce scenario is the running of the scalar spectral index (see [27, 28]).

³This could be easily further generalized to a Galileon field [43], which has also been done for inflation (see, e.g., [44–48]).

⁴The studies of Refs. [20, 51, 53] have been carried out for models where the nonsingular bounce is attributed to a noncanonical scalar field. Loop quantum cosmology (LQC) provides an alternative class of nonsingular bouncing models that could suppress r during the bounce. In LQC, the amplitude of the suppression depends on the equation of state during the bounce; if it is close to zero, then the suppression is very strong (see

in that it leads to the production of large non-Gaussianities [51], a problem that might be generic to a large class of nonsingular bounces [56, 57]. Again, these large non-Gaussianities are excluded by current measurements from the CMB [58]. This leads to conjecture that single field matter bounce cosmology suffers from a no-go theorem [51], which states that one cannot satisfy the bound on r without violating the bounds on non-Gaussianities and vice versa.

There is another way to suppress the tensor-to-scalar ratio if the sound speed of the perturbations can be smaller than the speed of light during the matter-dominated contracting phase. For example, in the Λ CDM bounce scenario [54] (and its extension [59]; see the review [28]), if there exists a form of dark matter with a small sound speed that dominates the contracting phase when the scale-invariant power spectra are generated, then the tensor-to-scalar ratio is already suppressed proportionally to the sound speed. Therefore, this provides another motivation to study non-Gaussianities when the sound speed is small during the matter-dominated contracting phase. An immediate question is whether the no-go theorem still holds true in this case or whether it can be circumvented. In this work, we want to explore this possibility of having a k -essence scalar field that would mimic dust-like matter with a small sound speed at low energies and that could play the role of the NEC violating scalar field during the bounce.

In this paper, we will evaluate the bispectrum produced by a k -essence scalar field in a matter-dominated contracting universe. This more general setup will yield richer features, which have the potential to be detected by future non-Gaussianity observations. In particular, the shapes, amplitudes, and scaling behaviors will be studied systematically. We will show that a small sound speed implies a large amplitude associated with the three-point function. Accordingly, we will claim that the no-go theorem is not circumvented but rather extended: in single field matter bounce cosmology, one cannot suppress the tensor-to-scalar ratio, either from the onset of the initial conditions in the matter contracting phase or from the amplification of the curvature perturbations during the bouncing phase, without producing large non-Gaussianities.

The outline of the paper is as follows. In section 2, we first introduce the background dynamics of the matter bounce scenario and introduce the class of k -essence scalar field models that we study in this paper. In section 3, we calculate the power spectra of curvature perturbations and tensor modes and show how a small sound speed coming from the k -essence scalar field allows for the suppression of the tensor-to-scalar ratio. We then consider the primordial non-Gaussianity in section 4. Using the in-in formalism, we evaluate every contribution to the three-point function and give a detailed analysis of the size and shapes of the resulting bispectrum. In section 5, we compute the amplitude parameter of non-Gaussianities in different limits and finally combine these results with the bound on the sound speed from section 3 to show that the no-go theorem in matter bounce cosmology is extended. We summarize our results in section 6. Throughout this paper, we use the mostly plus metric convention, and we define the reduced Planck mass to be $M_{\text{Pl}} = (8\pi G_{\text{N}})^{-1/2}$, where G_{N} is Newton's gravitational constant.

2 Setup and background dynamics

The idea of the matter bounce scenario is to begin with a matter-dominated contracting phase. At the background level, this corresponds to having a scale factor as a function of [50, 54, 55] and references therein for a discussion of LQC effects in nonsingular bouncing cosmology).

physical time given by

$$a(t) = a_B \left(\frac{t - \tilde{t}_B}{t_B - \tilde{t}_B} \right)^{2/3}, \quad (2.1)$$

and the Hubble parameter follows,

$$H(t) = \frac{2}{3(t - \tilde{t}_B)}, \quad (2.2)$$

where t_B corresponds to the time of the beginning of the bounce phase and \tilde{t}_B corresponds to the time at which the singularity would occur if no new physics appeared at high energy scales. Accordingly, a_B is the value of the scale factor at t_B . In terms of the conformal time τ defined by $d\tau = a^{-1}dt$, the scale factor is given by

$$a(\tau) = a_B \left(\frac{\tau - \tilde{\tau}_B}{\tau_B - \tilde{\tau}_B} \right)^2, \quad (2.3)$$

where τ_B and $\tilde{\tau}_B$ are the conformal times corresponding to t_B and \tilde{t}_B . Throughout the rest of this paper, the scale factor is normalized such that $a_B = 1$.

One can define the usual “slow-roll” parameters of inflation by

$$\epsilon \equiv -\frac{\dot{H}}{H^2} = \frac{3}{2}(1 + w), \quad \eta \equiv \frac{\dot{\epsilon}}{H\epsilon}, \quad (2.4)$$

where a dot denotes a derivative with respect to physical time, and $w \equiv p/\rho$ is the equation of state parameter with p and ρ denoting pressure and energy density, respectively. In the case of the matter bounce, the matter contracting phase implies that pressure vanishes, which is to say that

$$w = 0, \quad \epsilon = \frac{3}{2}, \quad \eta = 0. \quad (2.5)$$

If the pressure does not vanish exactly but is still very small, i.e. $|p/\rho| \ll 1$, then the values for w , ϵ , and η in equation (2.5) are only valid as leading order approximations, and they will be time dependent rather than constant. In this paper, we will work in the limit where equation (2.5) is valid.

In the usual matter bounce scenario, one would introduce a canonical scalar field to drive the matter-dominated contracting phase and describe the cosmological fluctuations. In this paper, we aim for more generality and assume that the perturbations are introduced by a k -essence scale field ϕ with Lagrangian density of the form⁵

$$\mathcal{L}_\phi = P(X, \phi), \quad (2.6)$$

where $X \equiv -\partial_\mu \phi \partial^\mu \phi / 2$, and we assume that the scalar field is minimally coupled to gravity. The energy density and pressure of this scalar field are then given by

$$\rho = 2XP_{,X} - P, \quad p = P, \quad (2.7)$$

where a comma denotes a partial derivative, e.g. $P_{,X} \equiv \partial P / \partial X$. Thus, the Friedmann equations read

$$3M_{\text{Pl}}^2 H^2 = 2XP_{,X} - P, \quad M_{\text{Pl}}^2 \dot{H} = -XP_{,X}. \quad (2.8)$$

⁵For an introduction to such a Lagrangian in early universe cosmology with the derivation of the background equations of motion and the definition of the different parameters, see, e.g., [34, 35, 38, 39].

Since we want a matter-dominated contracting phase, the pressure of the scalar field should vanish (at least in average), and $\rho = 2XP_{,X} \propto a^{-3}$.

It is helpful to have one specific example where a k -essence field drives the matter contraction. Let us consider the following Lagrangian density:

$$\mathcal{L}_\phi = K(X) = \frac{1}{8}(X - c^2)^2. \quad (2.9)$$

This type of Lagrangian belongs to a subclass of k -essence models $P(X, \phi)$ where the kinetic terms $K(X)$ are separate from the potential terms $V(\phi)$, i.e. $P(X, \phi) = K(X) - V(\phi)$. Moreover, the above Lagrangian has vanishing potential. Then, the ghost condensate solution is given by $X = c^2$ and $\phi(t) = ct + \pi(t)$, with $\dot{\pi}(t) \ll c$. In this case, the background equations yield $p \simeq 0$ and $\rho \sim \dot{\pi} \propto a^{-3}$, which exactly corresponds to a matter-dominated universe. More details about this model can be found in [60]. We note that there should be also other forms of $P(X, \phi)$ that can drive a matter contraction, and remarkably, the analysis that follows in this paper is done in a *model-independent* way and does *not* rely on the specific model of equation (2.9).

The sound speed and another “slow-roll” parameter are defined by⁶

$$c_s^2 \equiv \frac{\partial p}{\partial \rho} = \frac{P_{,X}}{P_{,X} + 2XP_{,XX}}, \quad s \equiv \frac{\dot{c}_s}{c_s H}. \quad (2.10)$$

Calculations will be done for a general sound speed, but as we will argue, we will be interested in the small sound speed limit, which can be realized with the appropriate form for $P(X, \phi)$. For instance, the explicit example given by equation (2.9) yields $c_s \simeq \dot{\pi}/c \ll 1$. Furthermore, we will generally assume later that the sound speed remains nearly constant, which is to say that $|s| \ll 1$. We also define two other variables for later convenience,

$$\Sigma \equiv XP_{,X} + 2X^2P_{,XX} = \frac{M_{\text{Pl}}^2 H^2 \epsilon}{c_s^2}, \quad (2.11)$$

and

$$\lambda \equiv X^2P_{,XX} + \frac{2}{3}X^3P_{,XXX} = \frac{X}{3}\Sigma_{,X} - \frac{1}{3}\Sigma. \quad (2.12)$$

The ratio λ/Σ will be of particular interest in the following sections. For inflation, it depends on the specific realization of the general single field, such as DBI and k -inflation models. For the matter bounce scenario, it can be obtained in an approximately model-independent way. The detailed calculation is in Appendix A, where we find that the ratio λ/Σ can be expressed in terms of the sound speed, as shown by equation (A.20).

3 Mode functions and two-point correlation functions

We begin with an action of the form

$$S = \int d^4x \sqrt{-g} \left(\frac{1}{2}M_{\text{Pl}}^2 R + \mathcal{L}_\phi \right), \quad (3.1)$$

where g is the determinant of the metric tensor and R is the Ricci scalar. Importantly, we assume that the matter Lagrangian \mathcal{L}_ϕ has the general form of equation (2.6), but we do

⁶We assume that the cosmological perturbations will remain adiabatic throughout the matter-dominated contracting phase.

not restrict our attention to any specific model. By perturbing up to second order the above action, one finds⁷

$$S^{(2)} = \int d\tau d^3\vec{x} z^2 \left[\zeta'^2 - c_s^2 (\vec{\nabla}\zeta)^2 \right] , \quad (3.2)$$

where $\zeta(\tau, \vec{x})$ denotes the curvature perturbation in the comoving gauge, i.e. on slices where fluctuations of the scalar field vanish ($\delta\phi = 0$). Also, a prime represents a derivative with respect to conformal time, $\vec{\nabla} = \partial_i$ is the spatial gradient, and we define $z^2 \equiv 2\epsilon a^2 M_{\text{Pl}}^2 / c_s^2$. Transforming to Fourier space, the second-order perturbed action becomes

$$S^{(2)} = \int d\tau \int \frac{d^3\vec{k}}{(2\pi)^3} z^2 \left[\zeta'(\vec{k})\zeta'(-\vec{k}) - c_s^2 k^2 \zeta(\vec{k})\zeta(-\vec{k}) \right] , \quad (3.3)$$

where $k^2 \equiv \vec{k} \cdot \vec{k} = |\vec{k}|^2$. Upon quantization of the curvature perturbation, one has

$$\hat{\zeta}(\tau, \vec{k}) = \hat{a}_{\vec{k}}^\dagger u_{\vec{k}}(\tau) + \hat{a}_{-\vec{k}} u_{\vec{k}}^*(\tau) , \quad (3.4)$$

where the annihilation and creation operators satisfy the usual commutation relation $[\hat{a}_{\vec{k}}, \hat{a}_{\vec{k}'}^\dagger] = (2\pi)^3 \delta^{(3)}(\vec{k} - \vec{k}')$. The equation of motion of the mode function is then given by

$$v_k'' + \left(c_s^2 k^2 - \frac{z''}{z} \right) v_k = 0 , \quad (3.5)$$

where the mode function is rescaled as $v_k = z u_k$ (v_k is called the Mukhanov-Sasaki variable). Together with the commutation relation $[\hat{\zeta}(\vec{k}_1), \hat{\zeta}'(\vec{k}_2)] = (2\pi)^3 \delta^{(3)}(\vec{k}_1 + \vec{k}_2)$, one finds (see, e.g., [42])

$$u_k(\tau) = \frac{iA[1 - ic_s k(\tau - \tilde{\tau}_B)]}{2\sqrt{\epsilon c_s k^3}(\tau - \tilde{\tau}_B)^3} e^{ic_s k(\tau - \tilde{\tau}_B)} \quad (3.6)$$

$$u_k'(\tau) = \frac{iA}{2\sqrt{\epsilon c_s k^3}} \left(\frac{-3[1 - ic_s k(\tau - \tilde{\tau}_B)]}{(\tau - \tilde{\tau}_B)^4} + \frac{c_s^2 k^2}{(\tau - \tilde{\tau}_B)^2} \right) e^{ic_s k(\tau - \tilde{\tau}_B)} \quad (3.7)$$

to be the solution to the equation of motion (3.5) in the context of a matter-dominated contracting universe as described in the previous section. Here, A is a normalization constant that is determined by the quantum vacuum condition at Hubble radius crossing in the contracting phase, which is given by $A = (\tau_B - \tilde{\tau}_B)^2 / M_{\text{Pl}}$.

The general two-point correlation functions are given by

$$\langle \hat{\zeta}(\tau_1, \vec{k}_1) \hat{\zeta}(\tau_2, \vec{k}_2) \rangle = (2\pi)^3 \delta(\vec{k}_1 + \vec{k}_2) u_{\vec{k}_1}^*(\tau_1) u_{\vec{k}_1}(\tau_2) , \quad (3.8)$$

$$\langle \hat{\zeta}(\tau_1, \vec{k}_1) \hat{\zeta}'(\tau_2, \vec{k}_2) \rangle = (2\pi)^3 \delta(\vec{k}_1 + \vec{k}_2) u_{\vec{k}_1}^*(\tau_1) u_{\vec{k}_1}'(\tau_2) , \quad (3.9)$$

and in particular, the power spectrum, evaluated at the bounce point τ_B (well after Hubble radius exit), is given by

$$\langle \hat{\zeta}(\tau_B, \vec{k}) \hat{\zeta}(\tau_B, \vec{k}') \rangle = (2\pi)^3 \delta^{(3)}(\vec{k} + \vec{k}') \frac{2\pi^2}{k^3} \mathcal{P}_\zeta(\tau_B, k) , \quad (3.10)$$

⁷Again, see, e.g., [30, 35, 38, 39] for a derivation of the perturbation equations in k -essence early universe cosmology.

where

$$\mathcal{P}_\zeta(\tau_B, k) = \frac{A^2}{8\pi^2 \epsilon c_s (\tau_B - \tilde{\tau}_B)^6} = \frac{1}{12\pi^2 c_s M_{\text{Pl}}^2 (\tau_B - \tilde{\tau}_B)^2} . \quad (3.11)$$

The scale invariance of the power spectrum in matter bounce cosmology is thus explicit from the above.

The above focused only on the scalar perturbations, but as mentioned in the introduction, the matter bounce scenario also generates a scale-invariant power spectrum of tensor perturbations. Considering the transverse and traceless perturbations to the spatial metric, $\delta g_{ij} = a^2 h_{ij}$, which can be decomposed as

$$h_{ij}(\tau, \vec{x}) = h_+(\tau, \vec{x}) e_{ij}^+ + h_\times(\tau, \vec{x}) e_{ij}^\times \quad (3.12)$$

with two fixed polarization tensors e_{ij}^+ and e_{ij}^\times , the second-order perturbed action has contributions of the form

$$S^{(2)} \supset \frac{M_{\text{Pl}}^2}{4} \int d\tau d^3\vec{x} a^2 \left[h'^2 - (\vec{\nabla} h)^2 \right] \quad (3.13)$$

for each polarization state h_+ and h_\times . By normalizing each state as $\mu = a M_{\text{Pl}} h/2$, the second-order perturbed action is of canonical form (μ is the Mukhanov-Sasaki variable), and the resulting equation of motion for each state is

$$\mu_k'' + \left(k^2 - \frac{a''}{a} \right) \mu_k = 0 , \quad (3.14)$$

where the equation is written in Fourier space. Since $a \sim \tau^2$ in a matter-dominated contracting phase, one has $a''/a = 2/\tau^2$, and so, one expects a scale-invariant power spectrum just as in de Sitter space. The tensor power spectrum is given by

$$\mathcal{P}_t = 2\mathcal{P}_h = 2 \left(\frac{2}{a M_{\text{Pl}}} \right)^2 \frac{k^3}{2\pi^2} |\mu_k|^2 , \quad (3.15)$$

where the first factor of 2 accounts for the two polarizations $+$ and \times , and the factor $[2/(a M_{\text{Pl}})]^2$ comes from the normalization of μ . Upon matching with quantum vacuum initial conditions at Hubble radius crossing similar to the above treatment for scalar modes, one finds the power spectrum of tensor modes at the bounce point to be given by

$$\mathcal{P}_t(\tau_B, k) = \frac{2}{\pi^2 M_{\text{Pl}}^2 (\tau_B - \tilde{\tau}_B)^2} , \quad (3.16)$$

which is indeed independent of scale.

The tensor-to-scalar ratio is then defined to be

$$r \equiv \frac{\mathcal{P}_t}{\mathcal{P}_\zeta} . \quad (3.17)$$

It follows from equations (3.11) and (3.16) that

$$r = 24 c_s \quad (3.18)$$

in the context of matter bounce cosmology with a general k -essence scalar field⁸. On one hand, this highlights the problem of standard matter bounce cosmology, which is driven by a

⁸Of course, this assumes that the perturbations remain constant on super-Hubble scales after the matter contraction phase, in particular through the bounce and until the beginning of the radiation-dominated expanding phase of standard Big Bang cosmology.

canonical scalar field with $c_s = 1$, in which case $r = 24$. On the other hand, the above result provides a natural mechanism to suppress the tensor-to-scalar ratio provided the k -essence scalar field has an appropriately small sound speed. For example, satisfying the observational bound [52] $r < 0.07$ at 95% confidence imposes a bound on the sound speed of the order of

$$c_s \lesssim 0.0029 . \quad (3.19)$$

4 Non-Gaussianity

The previous section showed that a k -essence scalar field could yield a small tensor-to-scalar ratio in the context of the matter bounce scenario. This is done at the expense of having a small sound speed. In what follows, the goal is to compute the bispectrum and see how a small sound speed affects the results.

4.1 Cubic action

To evaluate the three-point correlation function, we must expand the action (3.1) up to third order. Let us recall the result of [38], the third-order interaction action of a general single scalar field⁹,

$$\begin{aligned} S^{(3)} = \int dt d^3\vec{x} \left\{ -a^3 \left[\Sigma \left(1 - \frac{1}{c_s^2} \right) + 2\lambda \right] \frac{\dot{\zeta}^3}{H^3} + \frac{a^3 \epsilon}{c_s^4} (\epsilon - 3 + 3c_s^2) \zeta \dot{\zeta}^2 \right. \\ + \frac{a\epsilon}{c_s^2} (\epsilon - 2s + 1 - c_s^2) \zeta (\partial\zeta)^2 - 2a \frac{\epsilon}{c_s^2} \dot{\zeta} (\partial\zeta) (\partial\chi) + \frac{a^3 \epsilon}{2c_s^2} \frac{d}{dt} \left(\frac{\eta}{c_s^2} \right) \zeta^2 \dot{\zeta} \\ \left. + \frac{\epsilon}{2a} (\partial\zeta) (\partial\chi) \partial^2 \chi + \frac{\epsilon}{4a} (\partial^2 \zeta) (\partial\chi)^2 + 2f(\zeta) \frac{\delta L}{\delta \zeta} \Big|_1 \right\} , \end{aligned} \quad (4.1)$$

where it is understood that $(\partial\zeta)^2 = \partial_i \zeta \partial^i \zeta$, $(\partial\zeta)(\partial\chi) = \partial_i \zeta \partial^i \chi$, $\partial^2 \zeta = \partial_i \partial^i \zeta$, and where we define χ such that $\partial^2 \chi = a^2 \epsilon \dot{\zeta}$. Also, we have

$$\frac{\delta L}{\delta \zeta} \Big|_1 = a \left(\frac{d\partial^2 \chi}{dt} + H \partial^2 \chi - \epsilon \partial^2 \zeta \right) , \quad (4.2)$$

$$\begin{aligned} f(\zeta) = \frac{\eta}{4c_s^2} \zeta^2 + \frac{1}{c_s^2 H} \zeta \dot{\zeta} + \frac{1}{4a^2 H^2} \{ -(\partial\zeta)(\partial\zeta) + \partial^{-2} [\partial_i \partial_j (\partial^i \zeta \partial^j \zeta)] \} \\ + \frac{1}{2a^2 H} \{ (\partial\zeta)(\partial\chi) - \partial^{-2} [\partial_i \partial_j (\partial^i \zeta \partial^j \chi)] \} , \end{aligned} \quad (4.3)$$

where ∂^{-2} is the inverse Laplacian.

The first and second terms in the last line of equation (4.1) can be reexpressed as

$$\frac{\epsilon}{2a} (\partial\zeta)(\partial\chi) \partial^2 \chi + \frac{\epsilon}{4a} (\partial^2 \zeta) (\partial\chi)^2 = -\frac{a^3 \epsilon^3}{2} \zeta \dot{\zeta}^2 + \frac{\epsilon}{2a} \zeta (\partial_i \partial_j \chi) (\partial^i \partial^j \chi) + \mathcal{K} , \quad (4.4)$$

where the boundary term is given by

$$\mathcal{K} = \partial_i \left[\zeta (\partial^i \chi) (\partial^2 \chi) + \frac{1}{2} (\partial^i \zeta) (\partial\chi)^2 - \zeta (\partial^i \partial^j \chi) (\partial_j \chi) \right] . \quad (4.5)$$

⁹From here on, we take $M_{\text{Pl}} = 1$ for convenience.

Since the $\partial_i[\dots]$ term above does not contribute to the three-point function, the third-order action, equation (4.1), is equivalent to

$$\begin{aligned}
S^{(3)} = \int dt d^3\vec{x} \left\{ -a^3 \left[\Sigma \left(1 - \frac{1}{c_s^2} \right) + 2\lambda \right] \frac{\dot{\zeta}^3}{H^3} + \frac{a^3 \epsilon}{c_s^4} (\epsilon - 3 + 3c_s^2) \zeta \dot{\zeta}^2 \right. \\
+ \frac{a\epsilon}{c_s^2} (\epsilon - 2s + 1 - c_s^2) \zeta (\partial\zeta)^2 - 2a \frac{\epsilon}{c_s^2} \dot{\zeta} (\partial\zeta) (\partial\chi) + \frac{a^3 \epsilon}{2c_s^2} \frac{d}{dt} \left(\frac{\eta}{c_s^2} \right) \zeta^2 \dot{\zeta} \\
\left. - \frac{a^3 \epsilon^3}{2} \zeta \dot{\zeta}^2 + \frac{\epsilon}{2a} \zeta (\partial_i \partial_j \chi) (\partial^i \partial^j \chi) + 2f(\zeta) \frac{\delta L}{\delta \zeta} \Big|_1 \right\} .
\end{aligned} \tag{4.6}$$

In the case of a canonical field with $c_s = 1$, this action returns to equation (15) of [42]. Meanwhile, as usual the last term in this action is removed by performing the field redefinition

$$\zeta \rightarrow \tilde{\zeta} + f(\tilde{\zeta}) , \tag{4.7}$$

where $\tilde{\zeta}$ denotes the field after redefinition.

4.2 Contributions to the shape function

In this section, we calculate the three-point correlation function using the in-in formalism (to leading order in perturbation theory; see, e.g., [29–31] for the methodology),

$$\langle O(t) \rangle = -2 \operatorname{Im} \int_{-\infty}^t d\bar{t} \langle 0 | O(t) L_{\text{int}}(\bar{t}) | 0 \rangle , \tag{4.8}$$

where O represents a set of operators of the form $\hat{\zeta}^3$ in our case of interest. Then, the shape function, \mathcal{A} , is defined such that¹⁰

$$\langle \zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}_3} \rangle = (2\pi)^7 \delta^{(3)} \left(\sum_i \vec{k}_i \right) \frac{\mathcal{P}_\zeta^2}{\prod_i k_i^3} \mathcal{A}(\vec{k}_1, \vec{k}_2, \vec{k}_3) . \tag{4.9}$$

In what follows, we list all the contributions to the shape function coming from the field redefinition and the interaction action (4.6). It is easy to check that, when taking the limit $c_s = 1$, one recovers the results of [42] for the matter bounce with a canonical scalar field as expected.

4.2.1 Contribution from the field redefinition

In momentum space, the field redefinition can be written as

$$\zeta_{\vec{k}} \rightarrow \tilde{\zeta}_{\vec{k}} + \int \frac{d^3 \vec{k}_1}{(2\pi)^3} \left[-\frac{3}{2c_s^2} - \frac{3\epsilon}{4} \left(\frac{\vec{k}_1 \cdot (\vec{k} - \vec{k}_1)}{k_1^2} - \frac{(\vec{k} \cdot \vec{k}_1)[\vec{k} \cdot (\vec{k} - \vec{k}_1)]}{k^2 k_1^2} \right) \right] \tilde{\zeta}_{\vec{k}_1} \tilde{\zeta}_{\vec{k} - \vec{k}_1} . \tag{4.10}$$

This redefinition has the following contribution to the three-point correlation function,

$$\begin{aligned}
\langle \zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}_3} \rangle_{\text{redef}} = \int \frac{d^3 \vec{k}'}{(2\pi)^3} \left[-\frac{3}{2c_s^2} - \frac{3\epsilon}{4} \left(\frac{\vec{k}' \cdot (\vec{k}_3 - \vec{k}')}{k'^2} - \frac{(\vec{k}_3 \cdot \vec{k}')(\vec{k}_3 \cdot [\vec{k}_3 - \vec{k}'])}{k_3^2 k'^2} \right) \right] \\
\times \left(\zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}'} \zeta_{\vec{k}_3 - \vec{k}'} \right) + (2 \text{ permutations}) ,
\end{aligned} \tag{4.11}$$

¹⁰We use $\zeta_{\vec{k}_i}$ to refer to $\hat{\zeta}(\tau, \vec{k}_i)$ to simplify the notation from here on.

and accordingly, the contribution to the shape function is

$$\mathcal{A}_{\text{redef}} = \left(\frac{3\epsilon}{16} - \frac{3}{4c_s^2} \right) \sum_i k_i^3 + \frac{3\epsilon}{64} \sum_{i \neq j} k_i k_j^2 - \frac{3\epsilon}{64 \prod_i k_i^2} \left(\sum_{i \neq j} k_i^7 k_j^2 + \sum_{i \neq j} k_i^6 k_j^3 - 2 \sum_{i \neq j} k_i^5 k_j^4 \right). \quad (4.12)$$

When $c_s^2 \ll 1$, this contribution is enhanced compared to the canonical case.

4.2.2 Contribution from the $\zeta \dot{\zeta}^2$ term

The term $\zeta \dot{\zeta}^2$ in equation (4.6) yields the following contribution to the bispectrum

$$\begin{aligned} \langle \zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}_3} \rangle_{\zeta \dot{\zeta}^2} = & -2 \times 2 \operatorname{Im} \int_{-\infty}^{\tau_B} d\bar{\tau} (2\pi)^3 \delta \left(\sum_i \vec{k}_i \right) a^2 \left[\frac{\epsilon}{c_s^4} (\epsilon - 3 + 3c_s^2) - \frac{\epsilon^3}{2} \right] \\ & \times u_{k_1}^*(\tau_B) u_{k_1}(\bar{\tau}) u_{k_2}^*(\tau_B) u_{k_2}'(\bar{\tau}) u_{k_3}^*(\tau_B) u_{k_3}'(\bar{\tau}) + (2 \text{ permutations}). \end{aligned} \quad (4.13)$$

To leading order in $c_s k_i (\tau_B - \bar{\tau}_B) \ll 1$, i.e. on scales larger than the sound Hubble radius¹¹, and recalling the solutions for u_k and u_k' [equations (3.6) and (3.7)], we get the following contribution to the shape function,

$$\mathcal{A}_{\zeta \dot{\zeta}^2} = -\frac{c_s^2}{8} \left[\frac{1}{c_s^4} (\epsilon - 3 + 3c_s^2) - \frac{\epsilon^2}{2} \right] \sum_i k_i^3. \quad (4.14)$$

Again, when $c_s^2 \ll 1$, this contribution is enhanced compared to the canonical case.

4.2.3 Contribution from the $\dot{\zeta} \partial \zeta \partial \chi$ term

A similar computation for this term yields the following contribution to the shape function

$$\mathcal{A}_{\dot{\zeta} \partial \zeta \partial \chi} = -\frac{\epsilon}{8} \sum_i k_i^3 + \frac{\epsilon}{8 \prod_i k_i^2} \left(\sum_{i \neq j} k_i^7 k_j^2 - \sum_{i \neq j} k_i^4 k_j^5 \right). \quad (4.15)$$

We note that this contribution is independent of c_s .

4.2.4 Contribution from the $\zeta (\partial_i \partial_j \chi)^2$ term

For this term, the contribution to the shape function is given by

$$\begin{aligned} \mathcal{A}_{\zeta (\partial_i \partial_j \chi)^2} = & -\frac{c_s^2 \epsilon^2}{32} \sum_i k_i^3 + \frac{c_s^2 \epsilon^2}{64} \sum_{i \neq j} k_i^2 k_j^2 \\ & + \frac{c_s^2 \epsilon^2}{64 \prod_i k_i^2} \left(\sum_i k_i^9 - \sum_{i \neq j} k_i^6 k_j^3 + 3 \sum_{i \neq j} k_i^5 k_j^4 - 3 \sum_{i \neq j} k_i^7 k_j^2 \right). \end{aligned} \quad (4.16)$$

When $c_s^2 \ll 1$, this contribution is suppressed compared to the canonical case.

¹¹This is also called the Jeans radius; see [54, 61] for an explicit definition of this scale and its role in matter bounce cosmology when $c_s \neq 1$.

4.2.5 Contribution from the $\dot{\zeta}^3$ term

The $\dot{\zeta}^3$ term is a new element in the Lagrangian caused by the nontrivial sound speed, which does not show up in the cubic action of canonical fields. Its contribution to the bispectrum is

$$\begin{aligned} \langle \zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}_3} \rangle_{\dot{\zeta}^3} = & -6 \times 2 \operatorname{Im} \int_{-\infty}^{\tau_B} d\bar{\tau} (2\pi)^3 \delta^{(3)} \left(\sum_i \vec{k}_i \right) \left(-\frac{aM_{\text{Pl}}^2 \epsilon}{H c_s^2} \right) \left(1 - \frac{1}{c_s^2} + 2 \frac{\lambda}{\Sigma} \right) \\ & \times u_{k_1}^*(\tau_B) u'_{k_1}(\bar{\tau}) u_{k_2}^*(\tau_B) u'_{k_2}(\bar{\tau}) u_{k_3}^*(\tau_B) u'_{k_3}(\bar{\tau}) , \end{aligned} \quad (4.17)$$

where we have used the expression for Σ , equation (2.11). Then the contribution to the shape function is expressed as

$$\mathcal{A}_{\dot{\zeta}^3} = -\frac{9}{2} \left(1 - \frac{1}{c_s^2} + 2 \frac{\lambda}{\Sigma} \right) \sum_i k_i^3 . \quad (4.18)$$

Since this is a new contribution compared to the canonical case, it vanishes for $c_s^2 = 1$. Indeed, when $c_s^2 = 1$, $\lambda/\Sigma \simeq (1 - c_s^2)/(6c_s^2) = 0$ (see equation (A.20) in Appendix A) and $1 - 1/c_s^2 = 0$. We note though that when $c_s^2 \ll 1$, this contribution is large.

4.2.6 Secondary contributions

The contribution from the term

$$\frac{a^3 \epsilon}{2c_s^2} \frac{d}{dt} \left(\frac{\eta}{c_s^2} \right) \zeta^2 \dot{\zeta}$$

in equation (4.6) is exactly zero since $\eta = 0$ during the matter contraction. We can also neglect the contribution from the term

$$\frac{a\epsilon}{c_s^2} (\epsilon - 2s + 1 - c_s^2) \zeta (\partial \zeta)^2$$

since the leading order term of the resulting bispectrum is proportional to $c_s^2 k_i^2 (\tau_B - \tilde{\tau}_B)^2$, which means that it is suppressed outside the sound Hubble radius.

The above results differ from the ones of general single field inflation. As pointed out in [42], two main reasons account for the different non-Gaussianities between matter bounce cosmology and inflation. First, here the “slow-roll” parameter ϵ is of order one rather than being close to zero, so the amplitudes are larger and the higher-order terms in ϵ are not suppressed. Second, curvature perturbations grow on super-Hubble scales in a matter-dominated contracting universe, and this behaviour manifests itself in the integral of equation (4.8), while for inflation, ζ usually remains constant after horizon-exit, so there is no such contribution.

In what follows, we summarize the above results and give a detailed analysis of the bispectrum. In particular, the differences with the canonical single field matter bounce scenario are discussed.

4.3 Summary of results

One can gather all the contributions above and get the total shape function,

$$\begin{aligned} \mathcal{A}_{\text{tot}} = & \left(-\frac{105}{32} + \frac{39}{16c_s^2} + \frac{9c_s^2}{128} \right) \sum_i k_i^3 + \frac{3}{256}(3c_s^2 + 6) \sum_{i \neq j} k_i^2 k_j + \frac{3}{256 \prod_i k_i^2} \\ & \times \left[3c_s^2 \sum_i k_i^9 + (10 - 9c_s^2) \sum_{i \neq j} k_i^7 k_j^2 - (3c_s^2 + 6) \sum_{i \neq j} k_i^6 k_j^3 + (9c_s^2 - 4) \sum_{i \neq j} k_i^5 k_j^4 \right], \end{aligned} \quad (4.19)$$

where we have used $\epsilon = 3/2$ and $\lambda/\Sigma = (1 - c_s^2)/6c_s^2$ for the matter contraction stage. Now the only free parameter in the total shape function is the sound speed c_s . In what follows, we shall discuss several interesting aspects of this result.

4.3.1 Amplitude

The size of non-Gaussianity is depicted by the dimensionless amplitude parameter

$$f_{\text{NL}}(\vec{k}_1, \vec{k}_2, \vec{k}_3) = \frac{10}{3} \frac{\mathcal{A}_{\text{tot}}(\vec{k}_1, \vec{k}_2, \vec{k}_3)}{\sum_i k_i^3}. \quad (4.20)$$

As one can see in equation (4.19), for most values of $c_s \in (0, 1]$, the first term dominates the total shape function, and roughly, f_{NL} becomes

$$f_{\text{NL}} \simeq -\frac{175}{16} + \frac{65}{8c_s^2} + \frac{15c_s^2}{64}, \quad (4.21)$$

which yields $f_{\text{NL}} < 0$ for $0.87 \lesssim c_s \leq 1$ and $f_{\text{NL}} > 0$ for $c_s \lesssim 0.87$. Thus, besides the negative amplitude in the canonical case [42], a small sound speed in matter bounce cosmology can produce a positive f_{NL} . In the next section, we shall further discuss its behaviour in different limits to confront observations.

4.3.2 Shape

The shape of non-Gaussianity is described by the dimensionless shape function

$$\mathcal{F}(k_1/k_3, k_2/k_3) = \frac{\mathcal{A}_{\text{tot}}}{k_1 k_2 k_3}. \quad (4.22)$$

Then, the first term in equation (4.19) gives exactly the form of the local shape. Thus, when the prefactor of the first term is nonvanishing ($c_s \not\approx 0.87$), the shape function is dominated by the local form, while the remaining terms just give some corrections. The total shape of non-Gaussianity is shown in the left panel of Figure 1, which looks very similar to the plots in [42] for the canonical matter bounce except that the amplitude is much larger here with c_s small.

At the same time, this result differs from the one of general single field inflation, where the equilateral form dominates the shape of non-Gaussianity for $c_s \ll 1$ [38]. This is mainly caused by the different generation mechanisms of non-Gaussianity in these two scenarios. For the matter bounce scenario, the growth of curvature perturbations after Hubble radius exit makes a significant contribution to the final bispectrum. Meanwhile, the local form is usually

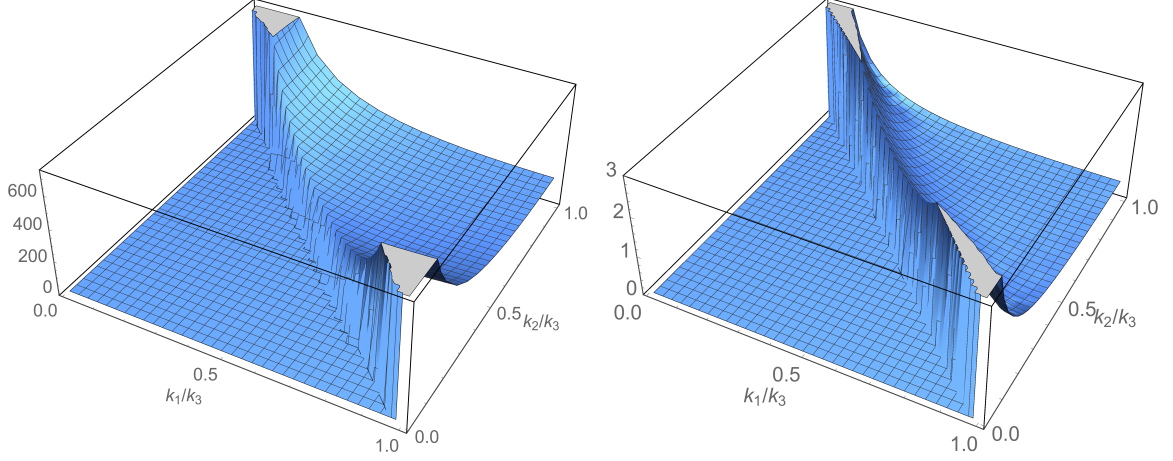


Figure 1. The shape of $\mathcal{F}(k_1/k_3, k_2/k_3)$ for $c_s = 0.2$ (left panel) and $c_s = 0.87$ (right panel).

thought to be generated on super-Hubble scales since “local” means that the non-Gaussianity at one place is *disconnected* with the one at other places. For general single field inflation, the dominant contribution is due to the enhanced interaction at horizon-crossing. Thus, these two scenarios behave quite differently with a small sound speed.

It is also interesting to note that for $c_s \approx 0.87$, the first term in equation (4.19) vanishes, so the shape function is dominated by the remaining terms. The shape of non-Gaussianity is plotted in the right panel of Figure 1 for this case, which is a new form different from the local one. To the best of our knowledge, no other scenario can give rise to such a kind of shape, thus it can be seen as a distinguishable signature of matter bounce cosmology for probes of non-Gaussianity.

4.3.3 The squeezed limit

Usually people are interested in the squeezed limit of the bispectrum ($k_1 \ll k_2 = k_3 = k$), since its scaling behaviour is helpful for clarifying the shapes of non-Gaussianity analytically. Here in the squeezed limit ($k_1/k \rightarrow 0$), the dimensionless shape function can be expanded as

$$\mathcal{F}(k_1/k_3, k_2/k_3) \simeq \frac{3}{8} \left(-\frac{33}{2} + \frac{13}{c_s^2} \right) \frac{k}{k_1} + \frac{3}{64} (1 + 6c_s^2) \frac{k_1}{k} + \mathcal{O} \left(\frac{k_1}{k} \right). \quad (4.23)$$

The leading order term gives the scaling $\mathcal{F} \sim k/k_1$ and

$$\langle \zeta_{\vec{k}_1} \zeta_{\vec{k}_2} \zeta_{\vec{k}_3} \rangle_{\text{squeezed}} \sim \frac{1}{k_1^3}, \quad (4.24)$$

which is consistent with the dominant local form. The only exception is when the coefficient of the first term vanishes ($c_s = \sqrt{26/33}$) and another scaling, $\mathcal{F} \sim k_1/k$, follows from the next-to-leading order term.

5 Amplitude parameter of non-Gaussianities and implication for the no-go theorem

There are three forms of the amplitude parameter f_{NL} that are of particular interest for cosmological observations. They are called the “local form”, the “equilateral form”, and the

“folded form”. The local form requires that one of the three momentum modes exits the Hubble radius much earlier than the other two, e.g., $k_1 \ll k_2 = k_3$. In this limit, one can simplify the total shape function, equation (4.19), to find

$$f_{\text{NL}}^{\text{local}} \simeq -\frac{165}{16} + \frac{65}{8c_s^2} . \quad (5.1)$$

The equilateral form requires that the three momenta form an equilateral triangle, i.e. $k_1 = k_2 = k_3$. In this case, we obtain

$$f_{\text{NL}}^{\text{equil}} \simeq -\frac{335}{32} + \frac{65}{8c_s^2} + \frac{45c_s^2}{128} . \quad (5.2)$$

The folded form has $k_1 = 2k_2 = 2k_3$, hence

$$f_{\text{NL}}^{\text{folded}} \simeq -\frac{37}{4} + \frac{65}{8c_s^2} . \quad (5.3)$$

As a result, in the limit where $c_s^2 \ll 1$, we find that

$$f_{\text{NL}}^{\text{local}} \approx f_{\text{NL}}^{\text{equil}} \approx f_{\text{NL}}^{\text{folded}} \approx \frac{65}{8c_s^2} \gg 1 . \quad (5.4)$$

Let us recall from section 3 that in order to satisfy the observational bound on the tensor-to-scalar ratio, we must impose $c_s \lesssim 0.0029$. This immediately implies

$$f_{\text{NL}}^{\text{local}} \approx f_{\text{NL}}^{\text{equil}} \approx f_{\text{NL}}^{\text{folded}} \gtrsim 9.55 \times 10^5 \gg 1 . \quad (5.5)$$

This amplitude of primordial non-Gaussianity is clearly ruled out according to the observations [58],

$$f_{\text{NL}}^{\text{local}} = 0.8 \pm 5.0 , \quad f_{\text{NL}}^{\text{equil}} = -4 \pm 43 , \quad f_{\text{NL}}^{\text{ortho}} = -26 \pm 21 , \quad (5.6)$$

thus ruling out the viability of the class of models studied here.

Alternatively, if one requires that, e.g., $-9.2 \lesssim f_{\text{NL}}^{\text{local}} \lesssim 10.8$ (i.e., imposing $f_{\text{NL}}^{\text{local}}$ to be within the measured 2σ error bars), then one would need¹² $c_s \gtrsim 0.62$. However, this lower bound on the sound speed yields a tensor-to-scalar ratio $r \gtrsim 14.88$, which is again clearly ruled out by observations [52].

In summary, there is no region of parameter space where c_s can give a good, small tensor-to-scalar ratio (i.e., of order 0.1 at most) and good, small non-Gaussianities (i.e., of order 10 at most). Therefore, independent of what happens during the bounce, we extend the no-go theorem conjectured in [51] to the following one:

No-Go Theorem. *For quantum fluctuations generated during a matter-dominated contracting phase, an upper bound on the tensor-to-scalar ratio (r) is equivalent to a lower bound on the amount of primordial non-Gaussianities (f_{NL}). Furthermore, if*

- *the matter contraction phase is due to a single (not necessarily canonical) scalar field,*
- *the same single scalar field allows for the violation of the NEC to produce a nonsingular bounce,*
- *and General Relativity holds at all energy scales,*

then satisfying the current observational upper bound on the tensor-to-scalar ratio cannot be done without contradicting the current observational upper bounds on f_{NL} (and vice versa).

¹²Note that this constraint does not exclude $c_s \approx 0.87$, for which the new shape of non-Gaussianity in the right panel of Figure 1 emerges.

6 Conclusions and discussion

In this paper, we computed the two- and three-point correlation functions produced by a generic k -essence scalar field in a matter-dominated contracting universe. Comparing the power spectra of scalar and tensor modes, we found that the tensor-to-scalar ratio can be appropriately suppressed if the sound speed associated with the k -essence scalar field is sufficiently small. In turn, we showed that the amplitude of the bispectrum is amplified by the smallness of the sound speed. As a result, it seems incompatible to suppress the tensor-to-scalar ratio below current observational bounds without producing excessive non-Gaussianities. This leads us to extend the conjecture of the no-go theorem, which effectively rules out a large class of nonsingular matter bounce models.

Although this seriously constrains nonsingular matter bounce cosmology as a viable alternative scenario to inflation, there remain several classes of models that are not affected by this no-go theorem. Indeed, one could still evade the no-go theorem assuming certain modified gravity models as stated in [51] (see references therein) or with the introduction of one or several new fields. For example, in the matter bounce curvaton scenario [62] (see also [15, 63, 64] for other nonsingular bouncing models using the curvaton mechanism) and in the two-field matter bounce scenario [65], entropy modes are generated by the presence of an additional scalar field, which are then converted to curvature perturbations. In both models near the bounce, the kinetic term of the entropy field varies rapidly, which acts as a tachyonic-like mass that amplifies (in a controlled way) the entropy fluctuations while not affecting the tensor modes. As a result, the tensor-to-scalar ratio is suppressed (see [10, 50] for reviews of this process). Furthermore, the production of non-Gaussianities in the matter bounce curvaton scenario has been estimated in [62], and it indicated that sizable, negative non-Gaussianities appeared, yet still in agreement with current observations. Accordingly, such a curvaton scenario does not appear to suffer from a no-go theorem. However, there still remains to do an appropriate extensive analysis of the production of non-Gaussianities when general multifields are included in the matter bounce scenario.

A similar curvaton mechanism is used in the new Ekpyrotic model [66, 67] (extended in [68–71]), which generates a nearly scale-invariant power spectrum of curvature perturbations. In this case, however, the smallness of the observed tensor-to-scalar ratio must be attributed to the fact that the tensor modes have a blue power spectrum when they exit the Hubble radius in a contracting phase with $w \gg 1$. The new Ekpyrotic model originally predicted large non-Gaussianities [72–76] (see also the reviews [77, 78]), but some more recent extensions can resolve this issue [79–83]. Thus, here as well, it appears that these types of models do not suffer from a similar no-go theorem¹³.

We note that one might be able to prove the no-go conjecture of this paper borrowing similar techniques to the effective field theory of inflation [40], i.e. by constructing an effective field theory of nonsingular bouncing cosmology (e.g., see the recent work of [25, 26]). In complete generality, this could allow us to find the exact and explicit relation between the tensor-to-scalar ratio (which involves the power spectra of curvature and tensor modes) and the bispectrum. In fact, the goal would be to find a consistency relation for the three-point function in single field nonsingular bouncing cosmology similar to what has been done in inflation [29, 85, 86]. This will be explored in a follow-up study.

¹³Furthermore, Ekpyrotic models are robust against the growth of anisotropies in a contracting universe. This is another challenge with the matter bounce scenario (see [18, 84]) that will have to be overcome to have a viable theory.

Finally, we would like to emphasize that, for matter bounce cosmology, although the simplest k -essence model is ruled out by the no-go theorem, the bispectrum with $c_s \neq 1$ (as an independent result of this paper) remains to be a probable target for future probes of non-Gaussianity. This possibility relies on the aforementioned bouncing models that can evade the no-go theorem with other mechanisms. In those cases, a nontrivial sound speed may still lead to the same behaviour of non-Gaussianities found in this paper, which potentially can be detected by future observations. Particularly, we predict a new shape with an amplitude still consistent with current observational limits, which can serve as the distinctive signature of matter bounce cosmology and help us distinguish it from other very early universe theories.

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A The ratio λ/Σ

Let us recall the definition of Σ and λ in equations (2.11) and (2.12). Their ratio is thus given by

$$\frac{\lambda}{\Sigma} = \frac{1}{3} \left(X \frac{\Sigma_{,X}}{\Sigma} - 1 \right). \quad (\text{A.1})$$

Recalling the definition of c_s^2 in equation (2.10), we note that

$$\Sigma = X(P_{,X} + 2XP_{,XX}) = X \frac{P_{,X}}{c_s^2}. \quad (\text{A.2})$$

Also, recalling the expression for ρ and p in equation (2.7), we find that $2XP_{,X} = \rho + p$, and so, the above expression for Σ becomes

$$\Sigma = \frac{\rho + p}{2c_s^2}. \quad (\text{A.3})$$

Consequently,

$$X \frac{\Sigma_{,X}}{\Sigma} = X \frac{\rho_{,X} + p_{,X}}{\rho + p} - 2X \frac{c_{s,X}}{c_s}. \quad (\text{A.4})$$

Working in the limit where $p = 0$, we note that $\rho = 2XP_{,X}$, and so, $p_{,X} = P_{,X} = \rho/(2X)$, which implies that $p_{,X}/\rho = 1/(2X)$. Also, $\rho_{,X} = p_{,X}/c_s^2$ from the definition of the sound speed, and thus,

$$\frac{\rho_{,X}}{\rho} = \frac{p_{,X}}{\rho c_s^2} = \frac{1}{2c_s^2 X}. \quad (\text{A.5})$$

Therefore, equation (A.4) in the limit where $p = 0$ becomes

$$X \frac{\Sigma_{,X}}{\Sigma} = \frac{1}{2c_s^2} + \frac{1}{2} - 2X \frac{c_{s,X}}{c_s} . \quad (\text{A.6})$$

Alternatively, one can evaluate the ratio λ/Σ as

$$\frac{\lambda}{\Sigma} = \frac{1}{3} \left(\frac{\Sigma_{,X}}{\Sigma} X - 1 \right) = \frac{1}{3} \left(\frac{\dot{\Sigma}}{\Sigma} \frac{X}{\dot{X}} - 1 \right) . \quad (\text{A.7})$$

Since we can write $\Sigma = H^2 M_{\text{Pl}}^2 \epsilon / c_s^2$ and recalling the definition of the slow-roll parameters in section 2, we get

$$\frac{\dot{\Sigma}}{H\Sigma} = -2\epsilon + \eta - 2s . \quad (\text{A.8})$$

Now, we note that we can write

$$\eta = \frac{\dot{\epsilon}}{H\epsilon} = \frac{\ddot{H}}{H\dot{H}} - 2 \frac{\dot{H}}{H^2} = \frac{\ddot{H}}{H\dot{H}} + 2\epsilon . \quad (\text{A.9})$$

Also, the Friedmann equation $M_{\text{Pl}}^2 \dot{H} = -X P_{,X}$ implies that

$$\frac{\ddot{H}}{H\dot{H}} = \frac{1}{H} \left(\frac{\dot{X}}{X} + \frac{\dot{P}_{,X}}{P_{,X}} \right) , \quad (\text{A.10})$$

and so,

$$\frac{\dot{X}}{HX} = \eta - 2\epsilon - \frac{\dot{P}_{,X}}{P_{,X}} . \quad (\text{A.11})$$

Therefore, combining equation (A.8) and the above yields

$$\frac{\dot{\Sigma}}{\Sigma} \frac{X}{\dot{X}} = \frac{-2\epsilon + \eta - 2s}{-2\epsilon + \eta - \frac{\dot{P}_{,X}}{P_{,X}}} . \quad (\text{A.12})$$

In the limit where $p = 0$, we recall that $\epsilon = 3/2$ and $\eta = 0$, and as a result,

$$\frac{\dot{\Sigma}}{\Sigma} \frac{X}{\dot{X}} = \frac{3 + 2s}{3 + \frac{\dot{P}_{,X}}{P_{,X}}} . \quad (\text{A.13})$$

Comparing the above with equation (A.6), since $(\dot{\Sigma}/\Sigma)(X/\dot{X}) = X\Sigma_{,X}/\Sigma$, we find

$$\frac{3 + 2s}{3 + \frac{\dot{P}_{,X}}{P_{,X}}} = \frac{1}{2c_s^2} + \frac{1}{2} - 2X \frac{c_{s,X}}{c_s} , \quad (\text{A.14})$$

but

$$-2X \frac{c_{s,X}}{c_s} = -2 \frac{X}{\dot{X}} \frac{\dot{c}_s}{c_s} = -2s \frac{HX}{\dot{X}} = \frac{-2s}{\eta - 2\epsilon - \frac{\dot{P}_{,X}}{P_{,X}}} , \quad (\text{A.15})$$

where the last equality follows from equation (A.11). Thus, equation (A.14), with $\epsilon = 3/2$ and $\eta = 0$, leaves us with

$$\frac{3}{3 + \frac{\dot{P}_{,X}}{P_{,X}}} = \frac{1}{2c_s^2} + \frac{1}{2} , \quad (\text{A.16})$$

and consequently,

$$\frac{\dot{P}_{,X}}{P_{,X}} = -3 \left(\frac{1 - c_s^2}{1 + c_s^2} \right) . \quad (\text{A.17})$$

As a result, equation (A.13) becomes

$$\frac{\dot{\Sigma}}{\Sigma} \frac{X}{\dot{X}} = X \frac{\Sigma_{,X}}{\Sigma} = \frac{1}{2c_s^2} \left(1 + \frac{2}{3}s \right) (1 + c_s^2) , \quad (\text{A.18})$$

and in the end, (A.7) is equivalent to

$$\frac{\lambda}{\Sigma} = \frac{1}{3} \left[\frac{1}{2c_s^2} \left(1 + \frac{2}{3}s \right) (1 + c_s^2) - 1 \right] . \quad (\text{A.19})$$

In the limit where $|s| \ll 1$, this reduces to

$$\frac{\lambda}{\Sigma} \simeq \frac{1}{3} \left[\frac{1 + c_s^2}{2c_s^2} - 1 \right] = \frac{1 - c_s^2}{6c_s^2} . \quad (\text{A.20})$$

In comparison, DBI inflation has $\lambda/\Sigma = (1 - c_s^2)/(2c_s^2)$ (see [38]).

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